

Homogeneous instantons in bigravity

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ABSTRACT: We study homogeneous gravitational instantons, conventionally called the Hawking-Moss (HM) instantons, in bigravity theory. The HM instantons describe the amplitude of quantum tunneling from a false vacuum to the true vacuum. Corrections to General Relativity (GR) are found in a closed form. Using the result, we discuss the following two issues: reduction to the de Rham-Gabadadze-Tolley (dRGT) massive gravity and the possibility of preference for a large e -folding number in the context of the Hartle-Hawking (HH) no-boundary proposal. In particular, concerning the dRGT limit, it is found that the tunneling through the so-called self-accelerating branch is exponentially suppressed relative to the normal branch, and the probability becomes zero in the dRGT limit. As far as HM instantons are concerned, this could imply that the reduction from bigravity to the dRGT massive gravity is ill-defined.

KEYWORDS: Solitons Monopoles and Instantons, Models of Quantum Gravity

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1 Introduction

The notion of a massive spin-2 graviton mediating the gravitational force has been the subject of much debate since the first proposal by Fierz and Pauli [1]. Among many issues, there was a fatal problem that it looked almost impossible to avoid a ghost in the scalar sector of the theory, called the Boulware-Deser (BD) ghost [2–4]. A breakthrough was firstly made by a non-linear construction of a ghost-free model by de Rham, Gabadadze and Tolley [5, 6], called the dRGT model, where in the decoupling limit, the BD ghost was removed by the introduction of a Minkowski reference metric (for a review, see [7, 8]). Soon after this success, the model was generalized to the full non-linear case [9, 10] and the absence of the BD ghost was proved for a generic but non-dynamic reference metric [11]. Then it was realized that a simple generalization of the non-dynamical reference metric to a dynamical one would lead to a non-linear bigravity theory without BD ghost [12, 13]. After that, a series of discoveries of the cosmological solutions and analysis of their corresponding perturbations have been done (see for example, [14]–[23] for dRGT model and [24]–[30] for bigravity).

At this stage, it is interesting to explore another cosmological application of the theory, namely quantum transitions between different vacua in the very early universe, particularly in the context of the cosmic landscape [33]. It may also shed light on the Cosmological Constant Problem (CCP) in the landscape of vacua [31]–[34].

Quantum transitions between vacua are described by instantons which are solutions of the field equations with the Euclidean signature. In the context of dRGT massive gravity,

the Hawking-Moss (HM) [35] and Coleman-De Luccia (CDL) [36] instantons were studied in [37, 38]. It was found that depending on the choice of the model parameters, the presence of a graviton mass may influence the tunneling rate, hence may affect the stability of a vacuum. One of the intriguing results from the analysis of the HM instanton is its effect on the Hartle-Hawking (HH) no-boundary wavefunction [39]: in contrast to GR where the HH no-boundary wavefunction exponentially disfavors a large number of e -folds necessary for successful inflation, the HH no-boundary wavefunction in dRGT massive gravity may have a peak at a sufficiently large value of the Hubble parameter for which one may obtain a sufficient number of e -folds of inflation [40].

However, in the dRGT model, one needs to introduce a non-dynamical, fiducial metric and fix it once and for all, which is rather unnatural. In particular, in the context of the cosmic landscape where a variety of geometries are probably realized, it is much more natural to render the fiducial metric dynamical [12]. In this paper, we investigate HM instantons, that is, homogeneous instantons in bigravity. We introduce two scalar fields which are minimally coupled respectively to the physical and fiducial metrics. We then construct a HM solution and evaluate its action. We find that there are two branches of solutions as in the dRGT case. For each branch we analyze the contribution from the interaction between the physical and fiducial metrics with special attention paid to the following two issues:

- (i). Reduction of the bigravity theory to the dRGT massive gravity. This is the limit $M_f^2/M_P^2 \rightarrow \infty$ where M_f is the Planck mass associated with the gravitational action of the fiducial metric. This limit is rather tricky because the total Euclidean action contains the HM action of the fiducial metric which is proportional to M_f^2 , which would diverge in the dRGT limit. We find that such a divergent term can be eliminated in one of the branches by a proper renormalization, while it cannot be eliminated in the other branch. As discussed in [41], this could imply that the dRGT massive gravity as the limit of bigravity is not necessarily well-defined.
- (ii). Possible preference to a large e -folding number for the Hartle-Hawking (HH) no-boundary wave function. We find that the bigravity model offers such a possibility. The same is true in the case of dRGT gravity [40], but it seems the bigravity model has more interesting cosmological implications because the probability depends also on the cosmological constant and the scalar potential in the fiducial side. Hence, a direct comparison to General Relativity (GR) implies that in the context of bigravity, there seems to be a much better chance to realize the consistency between the HH no-boundary proposal and the inflationary scenario.

It should be noted that in our model, we assume two matter sectors coupled to the physical and fiducial metrics, respectively. Hence, our model is free from a ghost mode, which differs from the case where the same matter sector couples to both metrics [55]–[60].

This paper is organized as follows. In section 2, we setup the Lagrangian for our model and formulate the equations of motion for homogeneous (HM) instantons. In section 3, we obtain HM solutions and study their implications. Section 4 is devoted to conclusion and future prospects.

Throughout the paper, the Lorentzian signature is set to be $(-, +, +, +)$.

2 Bigravity model

We consider a bigravity model with the following action [12]:

$$\begin{aligned}
S = & \frac{1}{2} \int d^4x \left[\sqrt{-g} M_{\text{P}}^2 (R_{\text{g}} - 2\lambda_{\text{g}}) + \sqrt{-f} M_{\text{f}}^2 (R_{\text{f}} - 2\lambda_{\text{f}}) \right] \\
& + m_{\text{g}}^2 M_{\text{e}}^2 \int d^4x \sqrt{-g} \sum_{n=1}^4 \alpha_n \mathcal{U}_n(\mathcal{K}) \\
& + \int d^4x \left[\sqrt{-g} \mathcal{L}_{\text{mg}} + \sqrt{-f} \mathcal{L}_{\text{mf}} \right],
\end{aligned} \tag{2.1}$$

where $g_{\mu\nu}$ is the physical metric, $f_{\mu\nu}$ is the fiducial metric, M_{P} and M_{f} are the Planck masses of the physical and fiducial metrics, respectively, and m_{g} is a coupling constant for the interactions between the two metrics with $\alpha_1, \alpha_2, \alpha_3$ and α_4 being arbitrary constants. For the remaining quantities, R is the Ricci scalar, λ is a cosmological constant, and $\mathcal{L}_{\text{m}}$ is the matter Lagrangian, and the subscripts g and f are attached for those of the physical and fiducial sectors, respectively. The mass M_{e} is defined by

$$M_{\text{e}} = (M_{\text{P}}^{-2} + M_{\text{f}}^{-2})^{-1/2}. \tag{2.2}$$

Thus in the dRGT massive gravity limit $M_{\text{f}}/M_{\text{P}} \rightarrow \infty$, m_{g} coincides with the Fierz-Pauli mass. As for the matter, to be specific, we focus on a minimally coupled scalar on each side,

$$\mathcal{L}_{\text{mg}} = -g^{\mu\nu} \partial_{\mu} \phi_{\text{g}} \partial_{\nu} \phi_{\text{g}} - V_{\text{g}}(\phi_{\text{g}}), \tag{2.3}$$

$$\mathcal{L}_{\text{mf}} = -f^{\mu\nu} \partial_{\mu} \phi_{\text{f}} \partial_{\nu} \phi_{\text{f}} - V_{\text{f}}(\phi_{\text{f}}). \tag{2.4}$$

The interaction terms in eq. (2.1) are defined as¹

$$\mathcal{U}_1(\mathcal{K}) = [\mathcal{K}] \equiv \mathcal{K}_{\mu}^{\mu}, \tag{2.5}$$

$$\mathcal{U}_2(\mathcal{K}) = \frac{1}{2!} \left([\mathcal{K}]^2 - [\mathcal{K}^2] \right), \tag{2.6}$$

$$\mathcal{U}_3(\mathcal{K}) = \frac{1}{3!} \left([\mathcal{K}]^3 - 3 [\mathcal{K}] [\mathcal{K}^2] + 2 [\mathcal{K}^3] \right), \tag{2.7}$$

$$\mathcal{U}_4(\mathcal{K}) = \frac{1}{4!} \left([\mathcal{K}]^4 - 6 [\mathcal{K}^2] [\mathcal{K}]^2 + 8 [\mathcal{K}^3] [\mathcal{K}] + 3 [\mathcal{K}^2]^2 - 6 [\mathcal{K}^4] \right), \tag{2.8}$$

where $\mathcal{K}_{\nu}^{\mu} = \delta_{\nu}^{\mu} - \left(\sqrt{g^{-1}f} \right)_{\nu}^{\mu}$.²

2.1 Euclidean action

By the Wick rotation $\tau = it$, the Euclidean version of the action (2.1) is obtained as $S_E = -iS[t = -i\tau]$.³ Correspondingly, in the semiclassical limit, the tunneling rate per unit time per unit volume is expressed in terms of the Euclidean action as

$$\Gamma/V = A e^{-B}; \quad B = S_E[\bar{g}_{\mu\nu, B}, \bar{\phi}_B] - S_E[\bar{g}_{\mu\nu, F}, \bar{\phi}_F], \tag{2.9}$$

¹We note that the action (2.1) can be equivalently written in a more compact way as shown in [12].

²It should be noted that the square root expression is defined by the relationship $\left(\sqrt{g^{-1}f} \right)_{\sigma}^{\mu} \left(\sqrt{g^{-1}f} \right)_{\nu}^{\sigma} = g^{\mu\sigma} f_{\sigma\nu}$.

³Here we note that the ‘Minkowski’ version of the instanton solutions have been studied in [17, 24, 25].

where $\{\bar{g}_{\mu\nu,B}, \bar{\phi}_B\}$ is the so-called bounce solution, or an instanton, a solution of the Euclidean equations of motion with appropriate boundary conditions, and $\{\bar{g}_{\mu\nu,F}, \bar{\phi}_F\}$ is the solution staying at the false vacuum [36]. Conventionally a bounce solution $\{\bar{g}_{\mu\nu,B}, \bar{\phi}_B\}$ is explored assuming $O(4)$ -symmetry, because it is often the case that an $O(4)$ -symmetric solution gives the lowest action for a wide class of scalar-field theories [42], hence dominates the tunneling process. It is therefore reasonable to assume the same even in the presence of gravity [36, 43–45].

Here we simply extend the above assumption to the Lorentzian-invariant bigravity theory by imposing the $O(4)$ -symmetric ansatz for both the physical and fiducial metrics:

$$ds_{g,E}^2 = N^2(\tau)d\tau^2 + a^2(\tau)d\Omega_3^2, \quad (2.10)$$

$$ds_{f,E}^2 = N_f^2(\tau)d\tau^2 + b^2(\tau)d\Omega_3^2, \quad (2.11)$$

where τ is the common Euclidean time parameter for both metrics and $d\Omega_3^2$ is the metric on a unit three sphere.

Inserting the ansatz (2.10) and (2.11) into the Euclidean version of the action (2.1), we obtain

$$\begin{aligned} S_E = 2\pi^2 \Big\{ & -3M_P^2 \int d\tau a \left(\frac{\dot{a}^2}{N} + N \right) - 3M_f^2 \int d\tau b \left(\frac{\dot{b}^2}{N_f} + N_f \right) \\ & + \int d\tau a^3 N \left[M_P^2 \lambda_g + V_g + \frac{\dot{\phi}_g^2}{2N^2} + m_g^2 M_e^2 \sum_{n=0}^3 A_n \left(\frac{b}{a} \right)^n \right] \\ & + \int d\tau b^3 N_f \left[M_f^2 \lambda_f + V_f + \frac{\dot{\phi}_f^2}{2N_f^2} + m_g^2 M_e^2 \sum_{n=0}^3 B_n \left(\frac{b}{a} \right)^{n-3} \right] \Big\}, \end{aligned} \quad (2.12)$$

where a dot means a derivative with respect to τ , and A_n and B_n are combinations of the parameters $\alpha_1, \alpha_2, \alpha_3$ and α_4 given by

$$A_0 = -4\alpha_1 - 6\alpha_2 - 4\alpha_3 - \alpha_4, \quad (2.13)$$

$$A_1 = 3B_0 = 3(\alpha_1 + 3\alpha_2 + 3\alpha_3 + \alpha_4), \quad (2.14)$$

$$A_2 = B_1 = -3(\alpha_2 + 2\alpha_3 + \alpha_4), \quad (2.15)$$

$$A_3 = \frac{1}{3}B_2 = \alpha_3 + \alpha_4, \quad (2.16)$$

$$B_3 = -\alpha_4. \quad (2.17)$$

2.2 Equations of motion

By varying the action (2.12) with respect to N and N_f , we obtain the ‘Friedmann’ equations,

$$\frac{\dot{a}^2}{N^2 a^2} = \frac{1}{a^2} - \frac{1}{3M_P^2} \left[-\frac{\dot{\phi}_g^2}{2N^2} + \Lambda_g(X, \phi_g) \right], \quad (2.18)$$

$$\frac{\dot{b}^2}{N_f^2 b^2} = \frac{1}{b^2} - \frac{1}{3M_f^2} \left[-\frac{\dot{\phi}_f^2}{2N_f^2} + \Lambda_f(X, \phi_f) \right], \quad (2.19)$$

where $\Lambda_g(X, \phi_g)$ and $\Lambda_f(X, \phi_f)$ are defined as

$$\Lambda_g(X, \phi_g) \equiv M_P^2 \lambda_g^{\text{eff}} + m_g^2 M_e^2 \sum_{n=0}^3 A_n X^n, \quad \lambda_g^{\text{eff}} \equiv \lambda_g + \frac{V_g(\phi_g)}{M_P^2}, \quad (2.20)$$

$$\Lambda_f(X, \phi_f) \equiv M_f^2 \lambda_f^{\text{eff}} + m_g^2 M_e^2 \sum_{n=0}^3 B_n X^{n-3}, \quad \lambda_f^{\text{eff}} \equiv \lambda_f + \frac{V_f(\phi_f)}{M_f^2}, \quad (2.21)$$

and $X \equiv b/a$.

We note that by inserting eqs. (2.18) and (2.19) into (2.12), one obtains the on-shell action,

$$S_E = 4\pi^2 \int d\tau \left[aN \left(-3M_P^2 + a^2 \Lambda_g(X, \phi_g) \right) + bN_f \left(-3M_f^2 + b^2 \Lambda_f(X, \phi_f) \right) \right]. \quad (2.22)$$

It should be noted that the interaction between physical and fiducial metrics is encoded in Λ_g and Λ_f , even though the above action looks like the sum of two independent Einstein gravity actions.

In addition to eqs. (2.18) and (2.19), by varying with respect to $a(\tau)$ and $b(\tau)$, one obtains the second order differential equations,

$$\ddot{a} = \frac{\dot{a}\dot{N}}{N} - \frac{aN^2}{3M_P^2} \left(\frac{\dot{\phi}_g^2}{N^2} + \Lambda_g \right) + \frac{m_g^2 M_e^2 aN}{6M_P^2} \sum_{n=0}^3 \left(\frac{b}{a} \right)^n \left[nN A_n + (n-3) N_f B_n \right], \quad (2.23)$$

$$\ddot{b} = \frac{\dot{b}\dot{N}_f}{N_f} - \frac{bN_f^2}{3M_f^2} \left(\frac{\dot{\phi}_f^2}{N_f^2} + \Lambda_f \right) - \frac{m_g^2 M_e^2 bN_f}{6M_f^2} \sum_{n=0}^3 \left(\frac{b}{a} \right)^{n-3} \left[nN A_n + (n-3) N_f B_n \right]. \quad (2.24)$$

In the case of GR, the Friedmann equation corresponds to the Hamiltonian constraint which represents the time reparameterization invariance. Therefore the time derivative of it does not give a new, independent equation. In the current case, however, only one of eqs. (2.18) and (2.19) corresponds to the Hamiltonian constraint. Therefore, the time derivative of one of them gives a new equation which should be consistent with the above second order differential equations. Taking the time derivative of eq. (2.18), one obtains

$$\begin{aligned} \ddot{a} &= \frac{a^2 \dot{\phi}_g}{6\dot{a}M_P^2} \left(\ddot{\phi}_g + \frac{\dot{a}\dot{\phi}_g}{a} - \frac{M_P^2 \dot{\lambda}_g^{\text{eff}}}{\dot{\phi}_g} \right) - \frac{a\Lambda_g}{3M_P^2} - \frac{m_g^2 M_e^2 a}{6M_P^2} \left(\frac{\dot{b}}{\dot{a}} \frac{a}{b} - 1 \right) \sum_{n=0}^3 n A_n \left(\frac{b}{a} \right)^n \\ &= -\frac{a}{3M_P^2} \left(\dot{\phi}_g^2 + \Lambda_g \right) - \frac{m_g^2 M_e^2 a}{6M_P^2} \sum_{n=0}^3 \left(\frac{b}{a} \right)^n \left[\frac{\dot{b}}{\dot{a}} (3-n) B_n - n A_n \right], \end{aligned} \quad (2.25)$$

where in the first step, we used eq. (2.20) while in the second step, eqs. (2.13)–(2.17) are used. Comparing eq. (2.25) with (2.23), for consistency one finds the constraint equation,

$$\left(\frac{\dot{b}}{\dot{a}} - N_f \right) \sum_{n=0}^3 (3-n) B_n \left(\frac{b}{a} \right)^n = 0. \quad (2.26)$$

The above constraint equation implies the existence of two branches of solutions:

- Branch I

$$N_f = \frac{\dot{b}}{\dot{a}}, \quad (2.27)$$

- Branch II

$$\sum_{n=0}^3 (3-n) B_n \left(\frac{b}{a}\right)^n = 0. \quad (2.28)$$

In the following subsections, we discuss these two branches separately.

2.3 Branch I

In this branch, the lapse function for the fiducial metric is fixed as

$$N_f = \frac{\dot{b}}{\dot{a}}. \quad (2.29)$$

Combining eqs. (2.18) with (2.19) and using eq. (2.29), one obtains the equation,

$$X^2 \equiv \left(\frac{b}{a}\right)^2 = \frac{M_f^2}{M_P^2} \left(\frac{-\dot{\phi}_g^2 + 2\Lambda_g(X, \phi_g)}{-\frac{\dot{\phi}_f^2}{N_f^2} + 2\Lambda_f(X, \phi_f)} \right), \quad (2.30)$$

where we have set $N = 1$ for simplicity by using the time reparameterization invariance of the theory. Using eqs. (2.20) and (2.21), one can explicitly express eq. (2.30) as an equation containing a series of X up to 4th order in its power,

$$\begin{aligned} A_3 X^4 + X^3 \left\{ A_2 - \frac{M_P^2}{M_f^2} \left[B_3 + \frac{1}{m_g^2 M_e^2} \left(M_f^2 \lambda_f^{\text{eff}}(\phi_f) - \frac{\dot{\phi}_f^2}{2N_f^2} \right) \right] \right\} + X^2 \left(A_1 - 3A_3 \frac{M_P^2}{M_f^2} \right) \\ + X \left[A_0 - A_2 \frac{M_P^2}{M_f^2} + \frac{1}{m_g^2 M_e^2} \left(M_P^2 \lambda_g^{\text{eff}}(\phi_g) - \frac{\dot{\phi}_g^2}{2} \right) \right] - \frac{A_1}{3} \frac{M_P^2}{M_f^2} = 0. \end{aligned} \quad (2.31)$$

Generally, the above equation is not easy to solve. However, in the special case when the scalar fields ϕ_g and ϕ_f are slowly varying so that we have $\dot{\phi}_g^2 \ll \Lambda_g$ and $\dot{\phi}_f^2/N_f^2 \ll \Lambda_f$, we may ignore the kinetic terms and the equation reduces to an algebra equation for X ,

$$\begin{aligned} A_3 X^4 + X^3 \left[A_2 - \frac{M_P^2}{M_f^2} \left(B_3 + \lambda_f^{\text{eff}} \frac{M_f^2}{m_g^2 M_e^2} \right) \right] + X^2 \left(A_1 - 3A_3 \frac{M_P^2}{M_f^2} \right) \\ + X \left(A_0 - A_2 \frac{M_P^2}{M_f^2} + \lambda_g^{\text{eff}} \frac{M_P^2}{m_g^2 M_e^2} \right) - \frac{A_1}{3} \frac{M_P^2}{M_f^2} = 0. \end{aligned} \quad (2.32)$$

Solving this equation one obtains $X = X(\phi_g, \phi_f) = b/a$. Thus a solution in this branch exists provided that the above equation has a real, positive root. In particular, in the case of our interest where ϕ_g and ϕ_f are homogeneous, which is the case of our current interest, the above gives an exact solution for $X = b/a$.

Before closing this subsection, we mention a particular case of the model parameters. As discussed in the above, eq. (2.32) is in general an algebraic equation for X . However,

for a particular set of the parameters, all the coefficients of the powers of X may vanish identically. In this case $X = b/a$ becomes unconstrained. This implies there will be substantially more varieties of solutions, including those with non-compact or non-trivial topologies. Interestingly, it seems this corresponds to the partially massless bimetric theory [46, 47] where the mass coincides with the Higuchi bound [48]. Furthermore, it seems to be also related to the conformal gravity [49]. Detailed discussion on this case is beyond the scope of the present paper. We plan to study this case in depth in a forthcoming paper [61].

2.4 Branch II

In this branch, one obtains the algebra solution for X ,

$$\frac{b}{a} = X_{\pm}, \quad X_{\pm} \equiv \frac{\alpha_2 + 2\alpha_3 + \alpha_4 \pm \sqrt{\alpha_3 + \alpha_3^2 + \alpha_2(\alpha_3 - \alpha_4) - \alpha_1(\alpha_3 + \alpha_4)}}{\alpha_3 + \alpha_4}. \quad (2.33)$$

A solution in this branch exists when the model parameters are in the range such that one of $X_{\pm}$ is real and positive.

We note that this branch is analogous to the ‘self-accelerating’ branch in the dRGT model [15, 16]. In ref. [21], it was found that this branch in dRGT massive gravity model suffers from a ghost problem, hence is considered to be an unhealthy branch. However, in extended massive gravity theories, this problem may be relieved. Moreover, it is this branch which exhibits various interesting features, including the case for the Hartle-Hawking wave function in quantum cosmology where successful inflation may be possible in massive gravity, in contrast to the case of GR [40]. Hence we also consider the HM instantons in this branch in the following.

3 Compact instantons: Hawking-Moss instantons

In this section, we focus on the HM instantons, that is, compact and homogeneous instanton solutions [35]. For the HM instantons, the scalar fields are at local maxima of their potentials, respectively. Therefore, $\phi_g = \phi_{g,\text{HM}} = \text{constant}$, $\phi_f = \phi_{f,\text{HM}} = \text{constant}$, $V'_g(\phi_{g,\text{HM}}) = V'_f(\phi_{f,\text{HM}}) = 0$, and $V''_g(\phi_{g,\text{HM}}) < 0$ and $V''_f(\phi_{f,\text{HM}}) < 0$. From eqs. (2.18) and (2.19), an HM solution takes the form,

$$a(\tau) = \sqrt{\frac{3M_{\text{P}}^2}{\Lambda_{g,\text{HM}}}} \sin \left(\sqrt{\frac{\Lambda_{g,\text{HM}}}{3M_{\text{P}}^2}} \tau \right), \quad (3.1)$$

$$b(\tau) = \sqrt{\frac{3M_{\text{f}}^2}{\Lambda_{f,\text{HM}}}} \sin \left(\sqrt{\frac{\Lambda_{f,\text{HM}}}{3M_{\text{f}}^2}} f(\tau) \right), \quad (3.2)$$

where the function $f(\tau)$ is defined as $\dot{f}(\tau) \equiv N_{\text{f}}$, and $\Lambda_{g,\text{HM}}$ and $\Lambda_{f,\text{HM}}$ are the values of Λ_g and Λ_f at $\phi_g = \phi_{g,\text{HM}}$ and $\phi_f = \phi_{f,\text{HM}}$, respectively. It has been shown in the previous section that the HM solutions in both branches satisfy $X = b/a = \text{constant}$. Consequently, one finds the expression for $\dot{f}(\tau)$ as

$$\dot{f}(\tau) = \frac{X \cos \left(\sqrt{\frac{\Lambda_{g,\text{HM}}}{3M_{\text{P}}^2}} \tau \right)}{\sqrt{1 - \frac{X^2 M_{\text{P}}^2}{M_{\text{f}}^2} \frac{\Lambda_{f,\text{HM}}}{\Lambda_{g,\text{HM}}} \sin^2 \left(\sqrt{\frac{\Lambda_{g,\text{HM}}}{3M_{\text{P}}^2}} \tau \right)}}. \quad (3.3)$$

On the other hand, from eqs. (2.18) and (2.19), the time parameters τ and $f(\tau)$ can be expressed in terms of the scale factors as

$$d\tau = \frac{da}{\sqrt{1 - \frac{a^2 \Lambda_g(a)}{3M_P^2}}}, \quad df = \frac{db}{\sqrt{1 - \frac{b^2 \Lambda_f(b)}{3M_f^2}}}. \quad (3.4)$$

Inserting these into the on-shell action (2.22), the Euclidean action may be expressed as

$$S_E = -12\pi^2 \left[M_P^2 \int_0^{a_{\max}^2} \sqrt{1 - \frac{a^2 \Lambda_g(a)}{3M_P^2}} da^2 + M_f^2 \int_0^{b_{\max}^2} \sqrt{1 - \frac{b^2 \Lambda_f(b)}{3M_f^2}} db^2 \right], \quad (3.5)$$

where $a_{\max}$ and $b_{\max}$ are the scale factors at their maxima.⁴ In the following, we compute the on-shell Euclidean action for both branches.

3.1 Euclidean action in branch I

In Branch I, from the fact that $X = b/a$ is a constant, it is straightforward to obtain the following relation:

$$N_f = \dot{f} = \frac{\dot{b}}{\dot{a}} = \frac{b}{a} = X, \quad (3.6)$$

where the parameter X is found to be fixed as

$$X = \frac{M_f}{M_P} \sqrt{\frac{\Lambda_{g, \text{HM}}}{\Lambda_{f, \text{HM}}}}. \quad (3.7)$$

We note that this is consistent with eq. (2.30). It also implies that the bubble expansion in the fiducial metric side synchronizes with the one in the physical side,

$$\begin{cases} a(\tau) = \sqrt{\frac{3M_P^2}{\Lambda_{g, \text{HM}}}} \sin \left(\sqrt{\frac{\Lambda_{g, \text{HM}}}{3M_P^2}} \tau \right), \\ b(\tau) = \sqrt{\frac{3M_f^2}{\Lambda_{f, \text{HM}}}} \sin \left(\sqrt{\frac{\Lambda_{g, \text{HM}}}{3M_P^2}} \tau \right). \end{cases} \quad (3.8)$$

Hence, the on-shell action (3.5) is obtained as

$$S_{E, \text{HM}}^{\text{B-I}} = -24\pi^2 \left(\frac{M_P^4}{\Lambda_{g, \text{HM}}} + \frac{M_f^4}{\Lambda_{f, \text{HM}}} \right), \quad (3.9)$$

From eq. (3.9), it is obvious that in Branch I, the system looks exactly like two copies of general relativity. This is partly because ‘ $b/a = \text{constant}$ ’ implies that the interaction term between the physical and fiducial metrics becomes a constant and it mimics an effective cosmological constant on each side, and partly because the relation (3.6) makes both metrics synchronize with each other, as shown in eq. (3.8).

⁴In this equation each of the integrals is done from 0 to its maximum, that is, a half of the corresponding 4-sphere. Thus one should multiply it by a factor of two to obtain the total action. This explains the coefficient $-12\pi^2$ instead of $-6\pi^2$.

Now let us consider the dRGT massive gravity limit, $M_f/M_P \rightarrow \infty$, in this branch. For concreteness, we assume λ_f is fixed. Thus in the limit $M_f/M_P \rightarrow \infty$, X remains finite, and hence so does M_f/Λ_f . Then it is obvious that the second term in eq. (3.9) diverges in this limit. Thus one might worry if the corresponding tunneling probability (2.9) would diverge. However, we argue that this divergence term is not a physical disaster, but may be removed by an appropriate renormalization.

To see this, we recall eq. (2.21) where the expression for Λ_f is given. If we take the limit $M_f/M_P \rightarrow \infty$ while keeping M_f^2/Λ_f finite, it corresponds to the limit where the all the energy scales on the fiducial side is kept finite while the gravity there becomes infinitely heavy and decoupled. In this limit we have

$$\frac{M_f^2}{\Lambda_f} \rightarrow \frac{1}{\lambda_f} \quad \text{as} \quad \frac{M_f}{M_P} \rightarrow \infty, \quad (3.10)$$

where λ_f is a bare cosmological constant on the fiducial metric side. It follows that the second term in eq. (3.9) becomes

$$-24\pi^2 \frac{M_f^4}{\Lambda_{f,\text{HM}}} \rightarrow -24\pi^2 \frac{M_f^2}{\lambda_f}. \quad (3.11)$$

Thus the limiting value is given solely in terms of the parameters of the theory, namely, the gravitational and cosmological constants. This implies one can subtract this term universally independent of the solutions. Namely, we define the renormalized action as

$$S' = S + 24\pi^2 \frac{M_f^2}{\lambda_f}. \quad (3.12)$$

Then for the HM solution in this branch we have

$$S'_{\text{E, HM}}{}^{\text{B-I}} = -24\pi^2 \left(\frac{M_P^4}{\Lambda_{g,\text{HM}}} + \frac{M_f^4}{\Lambda_{f,\text{HM}}} - \frac{M_f^2}{\lambda_f} \right). \quad (3.13)$$

In the limit $M_f/M_P \rightarrow \infty$, this reduces to the expression in the dRGT model,

$$S'_{\text{E, HM}}{}^{\text{B-I}} \rightarrow S_{\text{E, HM, dRGT}}{}^{\text{B-I}} = -24\pi^2 \frac{M_P^4}{\Lambda_{g,\text{HM}}}. \quad (3.14)$$

3.2 Branch II

In Branch II, using the relation $b = X_{\pm}a$ where $X_{\pm}$ given by eq. (2.33), the on-shell action (3.5) is obtained as

$$S_{\text{E, HM}}^{\text{B-II}} = -24\pi^2 \left\{ \frac{M_P^4}{\Lambda_{g,\pm}} + \frac{M_f^4}{\Lambda_{f,\pm}} \left[1 - \left(1 - \frac{X_{\pm}^2 M_P^2 \Lambda_{f,\pm}}{M_f^2 \Lambda_{g,\pm}} \right)^{\frac{3}{2}} \right] \right\}, \quad (3.15)$$

where

$$\Lambda_{g,\pm} = M_P^2 \lambda_{g,\text{HM}}^{\text{eff}} + m_g^2 M_e^2 \sum_{n=0}^3 A_n X_{\pm}^n, \quad \Lambda_{f,\pm} = M_f^2 \lambda_{f,\text{HM}}^{\text{eff}} + m_g^2 M_e^2 \sum_{n=0}^3 B_n X_{\pm}^{n-3}. \quad (3.16)$$

As clear from the above expression, the correction term in this case has a different form from that in Branch I unless the square-root term on the right-hand side of it vanishes, which happens when $X_{\pm}$ takes the value of eq. (3.7). As we have seen in the previous subsection, this is the condition for the synchronization of both metrics. Thus unless the two metrics are synchronized, the form of the correction term in Branch II is different from that in Branch I.

In the massive gravity limit $M_f/M_P \rightarrow \infty$, the Euclidean action reduces to

$$S_{\text{E, HM}}^{\text{B-II}} \longrightarrow S_{\text{E, HM, dRGT}}^{\text{B-II}} - \frac{8\pi^2}{F^2} \frac{M_f^2}{M_P^2} \left[1 - (1 - \alpha_{\text{HM}}^2)^{\frac{3}{2}} \right], \quad (3.17)$$

where

$$S_{\text{E, HM, dRGT}}^{\text{B-II}} \equiv -\frac{8\pi^2 M_P^2}{H_{\text{HM}}^2} \left[1 - \frac{X_{\pm} Y_{\pm}}{6} \left(\frac{m_g}{H_{\text{HM}}} \right)^2 A(\alpha_{\text{HM}}) \right] \quad (3.18)$$

while $F \equiv \sqrt{\lambda_f^{\text{eff}}/3}$, $H_{\text{HM}} \equiv \sqrt{\Lambda_{g,\pm}/3M_P^2}$, $\alpha_{\text{HM}} \equiv X_{\pm}F/H_{\text{HM}}$ and $A(\alpha) \equiv \left[2 - \sqrt{1 - \alpha^2} (2 + \alpha^2) \right] / \alpha^4$. We see that the first term in eq. (3.17) exactly coincides with the result in the dRGT massive gravity (the detailed calculations are given in appendix). However, unlike the case of Branch I, the second divergent term contains the variable $\alpha_{\text{HM}} \propto H_{\text{HM}}^{-1}$ which depends on the solution. This implies that it is impossible to remove this divergence completely in the massive gravity limit.

If we remove the solution-independent universal divergence as in the case of Branch I, we obtain

$$S'_{\text{E, HM}}^{\text{B-II}} \rightarrow S_{\text{E, HM, dRGT}}^{\text{B-II}} + \frac{8\pi^2}{F^2} \frac{M_f^2}{M_P^2} (1 - \alpha_{\text{HM}}^2)^{\frac{3}{2}}. \quad (3.19)$$

Since the second term is positive definite and divergent, we conclude that the probability of tunneling through the Branch II solution is exponentially suppressed and vanishes in the dRGT massive gravity limit. This is consistent with [41] where it is shown that this class of bigravity solutions are lost in this limit.

3.3 Hartle-Hawking wave function

To determine a wave function of the universe in quantum cosmology, Hartle and Hawking proposed a boundary condition that the path integral should be done over compact metrics with Euclidean signature [50]. This is called the Hartle-Hawking (HH) no-boundary proposal. However, when applied to the inflationary universe, it predicts an exponentially small probability for a sufficiently large number of e -folds which is necessary for successful inflation.

Recently we found that this deficit of the HH wave function may be removed in dRGT massive gravity [40]. Namely, if the dRGT massive gravity is realized at very high energy scales, the correction term in the action may completely change the behavior of the HH wave function and a sufficiently large number of e -folds may be realized at high probability. Inspired by this success in dRGT massive gravity, in this subsection, we examine the same issue in our bigravity model by using the HM solutions we obtained in the previous subsections.

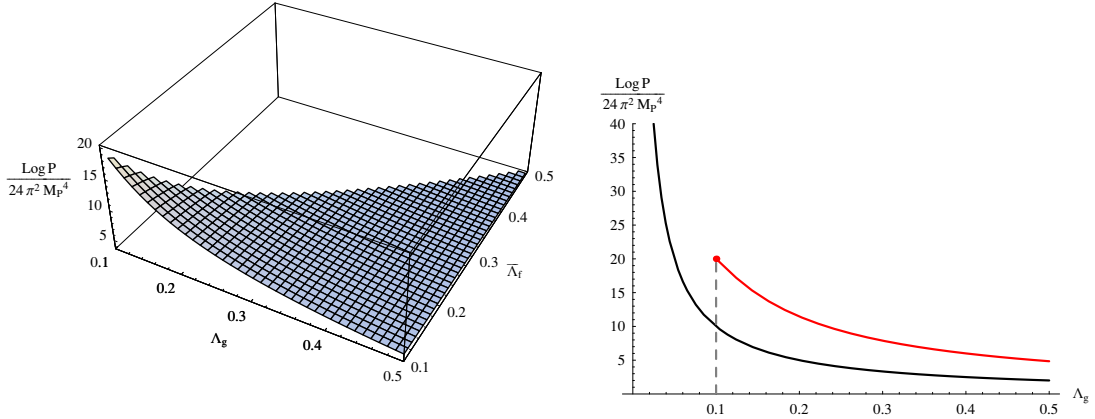


Figure 1. Left: the probability $\log P/24\pi^2 M_P^4$, with a fixed $\bar{X}_\pm = 1$. Right: for a given fixed $\bar{\Lambda}_{f,\pm} = 0.1$, the maximum probability appears around the cutoff (red), while there is no such cutoff for Einstein gravity (black).

The HH wave function is formally given by the path integral,

$$\Psi[h_{\mu\nu}, \phi_0] = \int^{\Sigma(h_{\mu\nu}, \phi_0)} \mathcal{D}g \mathcal{D}\phi e^{-S_E[g_{\mu\nu}, \phi]}, \quad (3.20)$$

where ϕ represents matter fields and the integration is over all regular and compact geometries M with the boundary $\partial M = \Sigma$ on which $\partial g = h$ and $\phi = \phi_0$. This may be extended in a straightforward manner to the case of bigravity by simply doubling the metric, $g \rightarrow (g, f)$ and $\partial M_g \oplus \partial M_f = \Sigma(h_g, h_f, \phi_0)$ where $(h_g, h_f) = (\partial g, \partial f)$. Here we focus on the mini-superspace and use the steepest-descent approximation to obtain

$$\Psi[a_0, b_0, \phi_0] = \int^{\Sigma(a_0, b_0, \phi_0)} \mathcal{D}a \mathcal{D}b \mathcal{D}\phi e^{-S_E[a, b, \phi]} \simeq \sum_{\text{sol}} e^{-S_E[a, b, \phi]}, \quad (3.21)$$

where the sum in the last term is over on-shell solutions. For a sufficiently flat potential the scalar field is slowly rolling, and one may approximate the scalar field to be a constant in time, $\phi \simeq \text{const}$ at leading order. In this case the HH wave function depends only on the value of ϕ through the effective cosmological constants for both physical and fiducial metrics, $\Lambda_g = \Lambda_g(\phi)$ and $\Lambda_f = \Lambda_f(\phi)$. The probability for a history that realizes $\phi = \phi_0$ is then given by

$$P(\phi_0) = |\Psi[\Lambda_g(\phi_0), \Lambda_f(\phi_0)]|^2 \propto e^{-2S_E[\Lambda_g(\phi_0), \Lambda_f(\phi_0)]}. \quad (3.22)$$

As we observed in eq. (3.9), for Branch I, the on-shell action is a simple sum of two independent actions, each of which has exactly the same form as that for Einstein gravity. Consequently the probability is dominated by the limit $\Lambda_{g,\text{HM}} \rightarrow 0$ as well as $\Lambda_{f,\text{HM}} \rightarrow 0$. On the other hand, for Branch II, the Euclidean action (3.15) gives the probability,

$$\frac{\log P}{24\pi^2 M_P^4} = \frac{1}{\Lambda_{g,\pm}} + \frac{1}{\bar{\Lambda}_{f,\pm}} \left[1 - \left(1 - \bar{X}_\pm^2 \frac{\bar{\Lambda}_{f,\pm}}{\Lambda_{g,\pm}} \right)^{3/2} \right], \quad (3.23)$$

where

$$\bar{\Lambda}_{f,\pm} \equiv \Lambda_{f,\pm} \frac{M_P^4}{M_f^4}, \quad \bar{X}_{\pm} \equiv \frac{M_P}{M_f} X_{\pm}. \quad (3.24)$$

Note that for Branch II, $X_{\pm}$ is given by the model parameters α_i , eq. (2.33). Hence the dependence of the probability on the solution is determined only by the values of Λ_f and Λ_g . As we can see from the left panel of figure 1, the allowed range of Λ_g is limited as

$$\Lambda_{g,\pm} \geq X_{\pm}^2 \frac{M_P^2}{M_f^2} \Lambda_{f,\pm}. \quad (3.25)$$

Hence, for a given $\Lambda_{f,\pm}$, the most probable value becomes $\Lambda_{g,\pm} = X_{\pm}^2 \Lambda_{f,\pm} M_P^2 / M_f^2$, as shown by the right panel of figure 1. The existence of this lower cutoff of Λ_g for a slowly rolling scalar field indicates that we may have an initial condition for inflation with a sufficiently large number of e -folds with sufficiently high probability.

4 Conclusion

As an approach to study non-perturbative effects in bigravity, we considered quantum tunneling by introducing two tunneling fields, respectively, minimally coupled to the physical and fiducial metrics. Then we derived for the Hawking-Moss (HM) instanton solutions. For a fixed set of the model parameters, we found two branches of solutions. We called these two branches as Branch I and II, respectively, and discussed their properties and implications.

First, we considered the dRGT massive gravity limit, $M_f/M_P \rightarrow \infty$, where the fiducial metric becomes non-dynamical. In this limit we found that the action diverges as $\propto M_f^2$, but in Branch I, the divergent term can be eliminated by a proper renormalization and the corresponding result in dRGT gravity is smoothly recovered. However, in Branch II, we found that the divergent term cannot be renormalized. Namely, there exists a solution-dependent divergence in the Euclidean action. This branch corresponds to the self-accelerating branch in the dRGT limit. Since this divergence is found to be positive definite, it implies that the probability of finding this branch is exponentially suppressed as we approach dRGT massive gravity. In dRGT massive gravity, the self-accelerating branch is known to be unstable due to the existence of a ghost mode [21]. Our result is quite interesting in this respect. It suggests that the self-accelerating branch may be avoided quantum cosmologically, if dRGT massive gravity is regarded as a limit in bigravity. On the other hand, this could also imply that the reduction from bigravity to the dRGT massive gravity may not be well-defined.

Second, as a direct application of the HM solution, we considered the wave function of the universe with the Hartle-Hawking (HH) no-boundary boundary condition. In this case, there is essentially no difference in the prediction of the Branch I solution from that of Einstein gravity. Namely, the HH wave function predicts the number of e -folds which is too small to make inflation successful. On the other hand, for the Branch II solution we found that the probability of realizing a sufficiently large number of e -folds becomes non-negligible, at least not exponentially suppressed. This suggests that the HH no boundary proposal may be saved in the context of bigravity.

It would be natural to go further to investigate the Coleman-De Luccia instantons in bigravity theory. In this case, the matter field is no more homogeneous and hence eq. (2.31) is no more an algebraic equation. That is, b/a is no longer a constant but varies as the scalar field varies. This makes the problem much more difficult to solve. We would like to come back to this topic in future.

Finally, as we mentioned at the end of section 2.3, for a particular case of the model parameters, $X = b/a$ becomes unconstrained. Hence this case will allow a lot more varieties of solutions, and may have intriguing cosmological implications [41]. Detailed discussion on this case is given in a forthcoming paper [61].

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A Massive gravity limit of branch II

In this appendix, we derive the dRGT massive gravity limit of the HM solutions given in eq. (3.15). For convenience, here we set $\alpha_1 = 0$ and $\alpha_1 = 1$ so that the value for $X_{\pm}$ coincides with that in the dRGT model. Since the first term in the curly brackets is nothing but the one for the conventional GR case, we focus on the reduction of the second term in the following.

In the limit $s \equiv M_P^2/M_f^2 \ll 1$, recalling that $M_e \rightarrow M_P$ from eq. (2.2), we have

$$\begin{aligned}
 & \frac{M_f^4}{\Lambda_{f,\pm}} \left[1 - \left(1 - \frac{X_{\pm}^2 M_P^2 \Lambda_{f,\pm}}{M_f^2 \Lambda_{g,\pm}} \right)^{\frac{3}{2}} \right] \\
 &= \frac{M_P^2}{s \left(\lambda_{f,\text{HM}}^{\text{eff}} + \frac{m_g^2 Y_{\pm}}{X^3} s \right)} \left\{ 1 - \left[1 - \frac{X_{\pm}^2 M_P^2}{\Lambda_{g,\pm}} \left(\lambda_{f,\text{HM}}^{\text{eff}} + \frac{m_g^2 Y_{\pm}}{X^3} s \right) \right]^{\frac{3}{2}} \right\} \\
 &= \frac{M_P^2}{s \lambda_{f,\text{HM}}^{\text{eff}}} \left(1 - \frac{m_g^2 Y}{X^3 \lambda_{f,\text{HM}}^{\text{eff}}} s + \mathcal{O}(s^2) \right) \left[1 - \gamma^{\frac{3}{2}} \left(1 - \frac{3\beta}{2} s \right) + \mathcal{O}(s^2) \right] \\
 &= \frac{M_P^2}{s \lambda_{f,\text{HM}}^{\text{eff}}} \left[1 - \gamma^{\frac{3}{2}} + s \left(\frac{3}{2} \gamma^{\frac{3}{2}} \beta + \frac{m_g^2 Y_{\pm}}{X^3 \lambda_{f,\text{HM}}^{\text{eff}}} (\gamma^{\frac{3}{2}} - 1) \right) + \mathcal{O}(s^2) \right], \quad (\text{A.1})
 \end{aligned}$$

where, for notational simplicity, we have introduced

$$\beta \equiv \frac{m_g^2 M_P^2 Y_{\pm}}{\gamma X_{\pm} \Lambda_{g,\pm}}, \quad \gamma \equiv 1 - \frac{M_P^2 X_{\pm}^2 \lambda_{f,\text{HM}}^{\text{eff}}}{\Lambda_{g,\pm}}, \quad (\text{A.2})$$

$$Y_{\pm} \equiv 3(1 - X_{\pm}) + 3\alpha_3(1 - X_{\pm})^2 + \alpha_4(1 - X_{\pm})^3 = X_{\pm}^3 \sum_{n=0}^3 B_n X_{\pm}^{n-3}. \quad (\text{A.3})$$

In order to compare the above with the result of dRGT massive gravity, we first note that from eq. (3.2), the scale factor $b(\tau)$ reduces as

$$b(\tau) \longrightarrow \sqrt{\frac{3}{\lambda_f^{\text{eff}}}} \sin \left(\sqrt{\frac{\lambda_f^{\text{eff}}}{3}} f(\tau) \right), \quad (\text{A.4})$$

while in dRGT massive gravity we have $b = F^{-1} \sin(Ff)$ with the fiducial Hubble parameter F . It is obvious that $\sqrt{\lambda_f^{\text{eff}}/3}$ plays the role of F ,⁵

$$\lambda_f^{\text{eff}} \longleftrightarrow 3F^2, \quad (\text{A.5})$$

while $\Lambda_{g,\pm}$ corresponds to the HM Hubble parameter,

$$\Lambda_{g,\pm} \longleftrightarrow 3M_{\text{P}}^2 H_{\text{HM}}^2. \quad (\text{A.6})$$

Hence, using eqs. (A.5) and (A.6), and from the definition of the dimensionless parameter $\alpha \equiv X_{\pm} F/H$, one finds the correspondences,

$$\frac{M_{\text{P}}^2 X_{\pm}^2 \lambda_f^{\text{eff}}}{\Lambda_{g,\pm}} \longleftrightarrow \alpha^2 \quad \implies \quad \gamma \longleftrightarrow 1 - \alpha^2. \quad (\text{A.7})$$

Inserting eqs. (A.5)–(A.7) into (A.1), one finds

$$\begin{aligned} & \frac{M_{\text{f}}^4}{\Lambda_{\text{f},\pm}} \left[1 - \left(1 - \frac{X_{\pm}^2 M_{\text{P}}^2 \Lambda_{\text{f},\pm}}{M_{\text{f}}^2 \Lambda_{g,\pm}} \right)^{\frac{3}{2}} \right] \\ & \longrightarrow \frac{M_{\text{P}}^2}{3F^2 s} \left[1 - (1 - \alpha^2)^{\frac{3}{2}} \right] + \frac{m_g^2 M_{\text{P}}^2 X_{\pm} Y_{\pm}}{18 H_{\text{HM}}^4 \alpha^4} \left[\sqrt{1 - \alpha^2} (2 + \alpha^2) - 2 \right]. \end{aligned} \quad (\text{A.8})$$

Inserting the above into eq. (3.15), one finally obtains the HM action in the dRGT massive gravity limit,

$$\begin{aligned} & S_{\text{E, HM}}^{\text{B-II}} \\ & \longrightarrow -\frac{8\pi^2 M_{\text{P}}^2}{H_{\text{HM}}^2} \left[1 - \frac{X_{\pm} Y_{\pm}}{6} \left(\frac{m_g}{H_{\text{HM}}} \right)^2 A(\alpha_{\text{HM}}) \right] - \frac{8\pi^2}{F^2} \frac{M_{\text{f}}^2}{M_{\text{P}}^2} \left[1 - (1 - \alpha_{\text{HM}}^2)^{\frac{3}{2}} \right], \end{aligned} \quad (\text{A.9})$$

where the function $A(\alpha)$ is defined by

$$A(\alpha) \equiv \frac{\left[2 - \sqrt{1 - \alpha^2} (2 + \alpha^2) \right]}{\alpha^4}. \quad (\text{A.10})$$

⁵It should be noted that in ref. [37], the Lorentzian signature of the fiducial metric is kept as it is throughout the computation since the fiducial metric is non-dynamical in dRGT massive gravity. However, it is dynamical in bigravity, so we should Wick rotate it as has been done in eq. (2.11). This introduces the appearance of an imaginary number in the function $f(\tau)$ which makes the hyperbolic function transform into the trigonometric function $b = F^{-1} \sin(Ff)$.

Comparing eq. (A.9) with eq. (4.14) of ref. [37], it is obvious that the first square brackets agrees with the result in dRGT massive gravity. However, unlike the case in Branch I, the second divergent term cannot be eliminated universally since it contains the solution-dependent variable $\alpha_{\text{HM}} \propto H_{\text{HM}}^{-1}$.

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